2010A Observing Proposals Due 30 September 2009; Survey Proposals Due 15 September 2009

Dave Bell

Standard proposals for NOAO-coordinated observing time for semester 2010A (February–July 2010) are **due by Wednesday evening, 30 September 2009, midnight MST**.

Proposals for new Survey programs are due by 15 September 2009, and require a letter of intent to have been sent in July. The facilities available this semester include the Gemini North and South telescopes, Cerro Tololo Inter-American Observatory (including SOAR), Kitt Peak National Observatory, and community-access time with Magellan, MMT, the Hale Telescope and the Center for High Angular Resolution Astronomy Array.

Proposal materials and information are available on our Web page (*www.noao.edu/noaoprop/*). There are four options for submission:

Web submissions—The Web form may be used to complete and submit all proposals. The information provided on the Web form is formatted and submitted as a LaTeX file, including figures that are "attached" to the Web proposal as encapsulated PostScript files.

File upload—A customized LaTeX file may be downloaded from the Web proposal form, after certain required fields have been completed. "Essay" sections can then be edited locally and the proposal submitted by uploading files through a Web page at: *www.noao.edu/noaoprop/submit/*.

Email submissions—A customized LaTeX file may be downloaded from the Web proposal form, after certain required fields have been completed. "Essay" sections can then be edited locally and the proposal submitted by email. Please carefully follow the instructions in the LaTeX template for submitting proposals and figures. Please use file upload instead of email if possible.

Gemini's Phase I Tool (PIT)—Investigators proposing for Gemini time **only** may optionally use Gemini's tool, which runs on Solaris, RedHat Linux, OS X, and Windows platforms and can be downloaded from *www.gemini.edu/sciops/P1help/p1Index.html*.

Note that proposals for Gemini time may also be submitted using the standard NOAO form, and that proposals that request time on Gemini plus other telescopes **MUST** use the standard NOAO form. PIT-submitted proposals will be converted for printing at NOAO, and are subject to the same page limits as other NOAO proposals. To ensure a smooth translation, please see the guidelines at: *www.noao.edu/noaoprop/help/pit.html*.

The addresses below are available to help with proposal preparation and submission:

Web proposal materials and information	*www.noao.edu/noaoprop/*
Request help for proposal preparation	*noaoprop-help@noao.edu*
Address for thesis and visitor instrument letters, as well as consent letters for use of PI instruments on the MMT	*noaoprop-letter@noao.edu*
Address for submitting LaTeX proposals by email	*noaoprop-submit@noao.edu*
Gemini-related questions about operations or instruments	*usgemini@noao.edu*
	www.noao.edu/gateway/gemini/support.html
CTIO-specific questions related to an observing run	*ctio@noao.edu*
KPNO-specific questions related to an observing run	*kpno@noao.edu*
MMT-specific questions related to an observing run	*mmt@noao.edu*
Magellan-specific questions related to an observing run	*magellan@noao.edu*
Hale-specific questions related to an observing run	*hale@noao.edu*

Community Access Time Available in 2010 with the CHARA Optical Interferometer Array

Steve Ridgway

NOAO and Georgia State University are announcing a one-time opportunity for observations with the Center for High Angular Resolution Astronomy (CHARA) Array at Mt. Wilson Observatory. About 50 hours will be available during calendar year 2010. Observations will be carried out by CHARA staff. This is intended primarily for scientists who would benefit from a small amount of data and wish to gain experience with optical interferometry capabilities.

Requests should be submitted using the standard NOAO proposal form by selecting "CHARA" in the telescope list, and entering "nights requested" as a decimal assuming 10 hours/night (e.g., 1.6 nights = 16 hours). Proposals must be submitted by the standard 2010A deadline of 30 September 2009. Note that this one-time call covers all of calendar year 2010, as opposed to the six-month period of February–July 2010 for other resources in the 2010A proposal cycle.

For more information, see *www.noao.edu/gateway/chara/*.

Community Access Time Available in 2010A with Magellan, MMT, and Hale

Dave Bell

As a result of awards made through the National Science Foundation's Telescope System Instrumentation Program (TSIP) and Renewing Small Telescopes for Astronomical Research (ReSTAR), telescope time is available to the general astronomical community at the following facilities in 2010A:

- **Magellan Telescopes**
  A total of five nights will be available for classically-scheduled observing programs with the 6.5-meter Baade and Clay telescopes at Las Campanas Observatory. For updated information on available instrumentation and proposal instructions, see *www.noao.edu/gateway/magellan/*.

- **MMT Observatory**
  Twelve nights of classically-scheduled observing time is to be available with the 6.5-meter telescope of the MMT Observatory. Previous requests have disproportionately used our allocation of dark and grey time, so bright time proposals are particularly encouraged. For further information, see *www.noao.edu/gateway/mmt/*.

- **Hale Telescope**
  Twelve nights of classically-scheduled observing time will be available with the 200-inch Hale Telescope at Palomar Observatory. Community access users may propose for the Double Spectrograph (a red and blue channel optical spectrograph) and TripleSpec (a near-infrared 1–2.5 μm spectrometer). For more information, see *www.noao.edu/gateway/hale/*.

General community access to the Keck telescopes will not be available in 2010A, but time will continue to be available with the HIRES spectrograph through a trade between Gemini and Keck. A list of instruments that we expect to be available in 2010A can be found at the end of this section. As always, investigators are encouraged to check the NOAO Web site for any last-minute changes before starting a proposal.

Observing Request Statistics for 2009B Standard Proposals

	No. of Requests	Nights Requested	Average Request	Nights Allocated	DD Nights (*)	Nights Previously Allocated	Nights Scheduled for New Programs	Over-subscription for New Programs
GEMINI								
GEM-N	154	151.01	0.98	59.55	0	0	59.55	2.54
GEM-S	93	108.42	1.17	56.4	0.5	0.25	56.15	1.93
CTIO								
CT-4m	48	174	3.62	139	0	9	130	1.34
SOAR	33	114.6	3.47	35.5	0	0.5	35	3.27
CT-1.5m	5	15.8	3.16	26.6	0	9.5	17.1	0.92
CT-1.3m	8	25.6	3.2	15.45	0	0	15.45	1.66
CT-1.0m	11	59	5.36	81	0	0	81	0.73
CT-0.9m	17	70.95	4.17	69.35	0	25	44.35	1.6
KPNO								
KP-4m	53	199	3.75	144	0	7	137	1.45
WIYN	30	94.8	3.16	60.5	0	10	50.5	1.88
KP-2.1m	40	231	5.78	148	0	0	148	1.56
KP-0.9m	4	15	3.75	12	0	0	12	1.25
Keck, MMT, Magellan								
Keck-I	33	52	1.58	8	0	0	8	6.5
Keck-II	33	42.1	1.28	7	0	0	7	6.01
Magellan-I	9	20	2.22	6	0	0	6	3.33
Magellan-II	6	13	2.17	4	0	0	4	3.25
MMT	8	13.92	1.74	13.5	0	0	13.5	1.03

*Nights allocated by NOAO Director

Gemini Instruments Expected to Be Available for 2010A

GEMINI NORTH	Detector	Spectral Range	Scale ("/pixel)	Field
NIRI	1024×1024 Aladdin Array	1–5μm R~500–1600	0.022, 0.050, 0.116	22.5", 51", 119"
NIRI + ALTAIR (AO- Natural or Laser)	1024×1024 Aladdin Array	1–2.5μm + L Band R~500–1600	0.022	22.5"
GMOS-N	3×2048×4608 CCDs	0.36–1.0μm R~670–4400	0.072	5.5' 5" IFU
Michelle	320×240 Si:As IBC	8–26μm R~100–30,000	0.10 img, 0.20 spec	32"×24" 43" slit length
NIFS	2048×2048 HAWAII-2RG	1–2.5μm R~5000	0.04×0.10	3"×3"
NIFS + ALTAIR (AO- Natural or Laser)	2048×2048 HAWAII-2RG	1–2.5μm R~5000	0.04×0.10	3"×3"

GEMINI SOUTH	Detector	Spectral Range	Scale ("/pixel)	Field
Phoenix	512×1024 Aladdin Array	1–5μm R<70,000	0.085	14" slit length
GMOS-S	3×2048×4608 CCDs	0.36–1.0μm R~670–4400	0.072	5.5' 5" IFU
T-ReCS	320×240 Si:As IBC	8–26μm R~100, 1000	0.09	28"×21"
NICI	1024×1024 (2 det.) Aladdin III InSb	0.9–5.5μm Narrowband Filters	0.018	18.4"×18.4"

EXCHANGE	Detector	Spectral Range	Scale ("/pixel)	Field
HIRES (Keck)	3×2048×4096 MIT-LL	0.35–1.0μm R~30,000–80,000	0.12	70" slit
MOIRCS (Subaru)	2×2048×2048 HAWAII-2	0.9–2.5μm R~500–3000	0.117	4'×7'
Suprime-Cam (Subaru)	10×2048×4096 CCDs	0.36–1.0μm	0.2	34'×27'
HDS (Subaru)	2×2048×4096 CCDs	0.3–1.0μm R~90,000	0.138	60" slit
FOCAS (Subaru)	2×2048×4096 CCDs	0.33–1.0μm R~250–7500	0.104	6' (circular)
COMICS (Subaru)	6×320×240 Si:As	8–25μm R~250, 2500, 8500	0.13	42"×32"
IRCS (Subaru)	1024×1024 InSb	1–5μm R~100–20,000	0.02, 0.05	21"×21", 54"×54"
IRCS+AO188 (Subaru)	1024×1024 InSb	1–5μm R~100–20,000	0.01, 0.02, 0.05	12×12", 21"×21", 54"×54"

CTIO Instruments Available for 2010A

Spectroscopy	Detector	Resolution	Slit
CTIO BLANCO 4m			
Hydra + Fiber Spectrograph	SITe 2K×4K CCD, 3300–11,000Å	700–18000, 45000	138 fibers, 2" aperture
R-C Spectrograph [1]	Loral 3K×1K CCD, 3100–11,000Å	300–5000	5.5'
SOAR 4.2m			
OSIRIS IR Imaging Spectrograph [2]	HgCdTe 1K×1K, JHK windows	1200, 1200, 3000	3.2', 0.5', 1.2'
Goodman Spectrograph [1,3]	Fairchild 4K×4K CCD, 3100–8500Å	1400, 2800, 6000	5.0'
CTIO/SMARTS 1.5m [4]			
Cass Spectrograph	Loral 1200×800 CCD, 3100–11,000Å	<1300	7.7'
Fiber echelle spectrograph	SITe 2K×2K CCD, 4020–7300Å	20000–42000	2.4" fiber

Imaging	Detector	Scale ("/pixel)	Field
CTIO BLANCO 4m			
Mosaic II Imager	8K×8K CCD Mosaic	0.27	36'
NEWFIRM [5]	InSb (mosaic, 4-2K×2K, 1–2.3μm)	0.4	28.0'
ISPI IR Imager	HgCdTe (2K×2K 1.0–2.4μm)	0.3	10.25'
SOAR 4.2m			
SOAR Optical Imager (SOI)	E2V 4K×4K Mosaic	0.08	5.25'
OSIRIS IR Imaging Spectrograph	HgCdTe 1K×1K	0.33, 0.14	3.2', 1.3'
Spartan IR Imager [6]	HgCdTe (mosaic 4-2K×2K)	0.068, 0.041	5.2', 3.1'
Goodman Spectrograph [3]	Fairchild 4K×4K CCD	0.15	7.2' diameter
CTIO/SMARTS 1.3m2 [7]			
ANDICAM Optical/IR Camera	Fairchild 2K×2K CCD	0.17	5.8'
	HgCdTe 1K×1K IR	0.11	2.0'
CTIO/SMARTS 1.0m [8]			
Direct Imaging	Fairchild 4K×4K CCD	0.29	20'
CTIO/SMARTS 0.9m [9]			
Direct Imaging	SITe 2K×2K CCD	0.4	13.6'

[1] The R-C Spectrograph should be out-performed by the Goodman Spectrograph on SOAR, in general. A comparison guide is available.

[2] The spectral resolutions and slit lengths for the OSIRIS imaging spectrograph correspond to its low-resolution, cross-dispersed, and high-resolution modes, respectively. In the cross-dispersed mode, one is able to obtain low-resolution spectra at JHK simultaneously.

[3] The Goodman Spectrograph is available in single-slit mode. Imaging mode is also available, but only with U, B, V, and R filters.

[4] Service observing only.

[5] Please see *www.noao.edu/ets/newfirm/* for more information. Permanently installed filters include J, H, and Ks. Please see NEWFIRM Web pages for updates on availability/schedulability of other filters.

[6] Some modes of the Spartan IR imager may be available. Please consult the NOAO Proposals Web pages for the latest information.

[7] Service observing only. Proposers who need the optical only will be considered for the 1.0-m unless they request otherwise. Note that data from both ANDICAM imagers is binned 2×2.

[8] Classical observing only. Observers may be asked to execute up to 1 hr per night of monitoring projects that have been transferred to this telescope from the 1.3-m. In this case, there will be a corresponding increase in the scheduled time. No specialty filters, no region of interest.

[9] Classical or service, alternating 7-night runs. If proposing for classical observing, requests for 7 nights are strongly preferred.

KPNO Instruments Available for 2010A

Spectroscopy	Detector	Resolution	Slit Length	Multi-object
Mayall 4m				
R-C CCD Spectrograph	T2KB/LB1A/F3KB CCD	300–5000	5.4'	single/multi
MARS Spectrograph	LB CCD (1980×800)	300–1500	5.4'	single/multi
Echelle Spectrograph	T2KB/F3KB CCD	18,000–65,000	2.0'	
FLAMINGOS[1]	HgCdTe (2048×2048, 0.9–2.5μm)	1000–1900	10.3'	single/multi
IRMOS[2]	HgCdTe (1024×1024, 0.9–2.5μm)	300/1000/3000	3.4'	single/multi
WIYN 3.5m				
Hydra + Bench Spectrograph[3]	STA1 CCD	700–22,000	NA	~85 fibers
SparsePak[4]	STA1 CCD	700–22,000	IFU	~82 fibers
2.1m				
GoldCam CCD Spectrograph	F3KA CCD	300–4500	5.2'	
FLAMINGOS[1]	HgCdTe (2048×2048, 0.9–2.5μm)	1000–1900	20.0'	

Imaging	Detector	Spectral Range	Scale ("/pixel)	Field
Mayall 4m				
CCD Mosaic-1	8K×8K	3500–9700 Å	0.26	35.4'
SQIID	InSb (3-512×512 illuminated)	JHKs	0.39	3.3'
FLAMINGOS[1]	HgCdTe (2048×2048)	JHK	0.32	10.3'
WIYN 3.5m				
Mini-Mosaic[5]	4K×4K CCD	3300–9700 Å	0.14	9.3'
OPTIC[5]	4K×4K CCD	3500–10,000 Å	0.14	9.3'
WHIRC[6]	VIRGO HgCdTe (2048×2048)	0.9–2.5μm	0.10	3.3'
2.1m				
CCD Imager[7]	T2KB CCD	3300–9700 Å	0.305	10.4'
SQIID	InSb (3-512×512 illuminated)	JHKs	0.68	5.8'
FLAMINGOS[1]	HgCdTe (2048×2048)	JHK	0.61	20.0'
WIYN 0.9m				
CCD Mosaic-1	8K×8K	3500–9700 Å	0.43	59'

[1] FLAMINGOS Spectral Resolution given assuming 2-pixel slit. Not all slits cover full field; check instrument manual. FLAMINGOS was built by the late Richard Elston and his collaborators at the University of Florida. Dr. Steve Eikenberry is currently the PI of the instrument.

[2] IRMOS, built by Dr. John MacKenty and collaborators. Availability will depend on proposal demand and block scheduling constraints.

[3] The Bench Spectrograph has recently been upgraded. A new CCD (STA 1), a new collimator, and two new Volume Phase Holographic (VPH) gratings, (740 l/mm and 3300 l/mm), are now available for use. Dispersion and wavelength range remain essentially the same in the upgraded system. However, observers should view www.wiyn.org/instrument/bench_upgrade.html for details on changes in throughput and instrumental resolution, as well as new options such as binning and gain choices, to help plan observations.

[4] Integral Field Unit, 80"×80" field, 5" fibers, graduated spacing.

[5] OPTIC Camera from University of Hawai'i is anticipated to be available at WIYN in 2010A through an agreement with Dr. John Tonry of the University of Hawai'i. At the present time, the exact time period of its availability is not defined. This instrument may be assigned to those that request to use Mini-Mosaic if this substitution still meets proposed imaging needs and making such an assignment would further observatory support constraints. Fast guiding mode of operation of OPTIC is now a supported mode for NOAO users of the instrument.

[6] WHIRC, built by Dr. Margaret Meixner (STScI) and collaborators, will be available for use during 2010A. Observers contemplating use with WTTM correction should consult with Dick Joyce or Lori Allen for details.

[7] While T2KB is the default CCD for CFIM, use of F3KB may be justified for some applications and may be specifically requested; scale 0.19"/pix, 9.7'×3.2' field. If T2KB is unavailable, CFIM may be offered with T2KA (scale 0.305"/pix, 10.4' field) or with F3KB to best match proposal requirements. www.noao.edu/kpno/ccdchar/ccdchar.html

MMT Instruments Available for 2010A

	Detector	Resolution	Spectral Range	Scale ("/pixel)	Field
BCHAN (spec, blue-channel)	Loral 3072×1024	R~800–11,000	0.32–0.8μm	0.3	150" slit
RCHAN (spec, red-channel)	Loral 1200×800	R~300–4000	0.5–1.0μm	0.3	150" slit
Hectospec (300-fiber MOS, PI)	2 2048×4608	R~1000–2000	0.38–1.1μm		60'
Hectochelle (240-fiber MOS, PI)	2 2048×4608	R~34,000	0.38–1.1μm		60'
ARIES (near-IR imager, PI)	1024×1024 HgCdTe		1.1–2.5μm	0.04, 0.02	20", 40"
CLIO (thermal-IR AI camera, PI)	320×256 InSb		H,K,L,M	0.05	16×13"
MIRAC3-BLINC (mid-IR img+nuller, PI)	128×128 Si:As BIB		2–25μm	0.09	11.5"
PISCES (wide n-IR imager, PI)	1024×1024 HgCdTe		1–2.5μm	0.18	3.1'
SWIRC (wide n-IR imager, PI)	2048×2048 HAWAII-2		1.0–1.6μm	0.15	5'
SPOL (img/spec polarimeter, PI)	Loral 1200×800	R~300–2000	0.38–0.9μm	0.2	20"

Magellan Instruments Available for 2010A

	Detector	Resolution	Spectral Range	Scale ("/pixel)	Field
Magellan I (Baade)					
PANIC (IR imager)	1024×1024 Hawaii		1–2.5μm	0.125	2'
IMACS (img/lslit/mslit)	8192×8192	R~2100–28,000	0.34–1.1μm	0.11, 0.2	15.5', 27.2'
MagIC (optical imager)	2048×2048		BVRI, u'g'r'i'z'	0.07	2.36'
Magellan II (Clay)					
LDSS3 (mslit spec/img)	4096×4096	R~200–1700	0.4–0.8μm	0.19	8.25' circ.
MIKE (echelle)	2K×4K	R~22,000	0.32–1.0μm	0.12–0.13	5" slit
MIKE Fibers (echelle)	2K×4K	R~16,000	0.32–1.0μm	0.12–0.13	20–23', 256 fibers
MagE (echellette)	1024×2048 E2V	R~4100	0.31–1.0μm	0.3	10" slit

Hale Instruments Available for 2010A

	Detector	Resolution	Spectral Range	Scale ("/pixel)	Field
Double Spectrograph/Polarimeter	1024×1024 red, 2048×4096 blue	R~1000–10,000	0.3–1.0μm	0.4–0.6	128" long, 8×15" multi
TripleSpec	1024×2048	R~2500–2700	1.0–2.4μm	0.37	30" slit

CHARA Instruments Available for 2010

	Beam Combiner	Resolution	Spectral Range	Beams
The CHARA Array consists of six 1-m aperture telescopes with baselines from 30 to 330 meters	Classic	Broadband	H or K	2
	MIRC	40	H or K	4
	Vega	1700	45 nm in V or R	2